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The period of $9^h 3^m 21^s.57$, assigned by PRACKA,¹ represents the observed maximum on July 17th exactly. Finally, it may be remarked that the photovisual amplitude is in good agreement with that found visually at the Laws Observatory.

F. H. SEARES and HARLOW SHAPLEY.

MT. WILSON SOLAR OBSERVATORY, August 18, 1914.

THE STAR CLUSTER N. G. C. 6760.

The object N. G. C. 6760 has been suspected as variable, but DREYER in a note in the N. G. C. states that there is no reason for supposing this to be so.

BORRELLY'S note in *C. R.*, **157**, 385, leads one to suppose that the object is changing in brightness, and it was accordingly placed on the observing program of the 60-inch reflector.

Two photographs were obtained on Seed 27 plates, with exposures of $\frac{3}{4}$ and $1\frac{1}{2}$ hours respectively. The $1\frac{1}{2}$ -hour photograph revealed a star cluster lying in a rich region of the Milky Way. A single count of this plate was made with a reseau having squares $30''$ on a side and the results plotted. It proved to be a cluster of 1,200 stars, $7'$ in diameter, lying on a background with a density of fifteen stars per square minute. The greatest number of stars counted in a single square of the reseau was 71, which would make 67 for the cluster. Mr. SEARES estimates that the stars in the cluster lie between 16.5 or 17 and 19.5 photographic magnitude.

No trace of nebulosity is found.

Under a power of 300, which was as much as the seeing would permit, the cluster appeared granular, but save for a few of the brighter stars was not certainly resolvable.

A comparison of the two plates under the stereocomparator gave no indication of variability.

FRANCIS G. PEASE.

MT. WILSON SOLAR OBSERVATORY.

SPECTRA OF STARS IN THE HERCULES CLUSTER M 13.

An additional photograph of Messier 13 was obtained with the focal plane spectrograph of the 60-inch reflector, using a slit width of 0.050^{mm} and an exposure of thirty hours.

¹ *V. J. S.*, 1913, p. 290.

The slit was set 3' south of the star Scheiner 373, which BARNARD uses as his normal star, and stars are scattered over a distance of 4' in an east and west line. Identification of twenty spectra is possible, though several are composite of several close stars. Spectra of some stars of Barnard visual magnitude 15 are included.

Taking as units five divisions of the Harvard scale, these stars are classified as follows:—

A₀ = 1
A₅ = 6
F₀ = 9
F₅ = 3
G₀ = 1

The mean of this determination is practically the same as that given by the nineteen stars mentioned in Volume XXV, page 260, of these *Publications*.

The two results combined, making a total of thirty-nine stars, give the mean class of the cluster as F₀.

FRANCIS G. PEASE.

MT. WILSON SOLAR OBSERVATORY.

NOTE ON THE BINARY STAR A 570.

In 1903 the star B. D. + 27°.2388 was found to be a close double. From the magnitude (5.90, Class A spectrum in *R. H. P.*) and the small angular separation, it was reasonably certain that the pair was a binary system, and the fact that the proper-motion was well defined, amounting to 0".088 in 236°.8 according to BOSS, confirmed this conclusion.

Subsequent measures have shown that the angular motion is rapid and it is probable that the period of revolution will be short, possibly in the neighborhood of twenty years.

The mean results of my measures are as follows:—

1903.40.....	198°.6	0".20	4 nights, 36-inch
1906.41.....	173 .0	0 .24	3 " "
1907.52.....	166 .8	0 .24	2 " "
1911.45.....	128 .3	0 .22	3 " "
1914.58.....	88 .5	0 .16	2 " "

R. G. AITKEN.

August, 1914.